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Stimulated radiative association of sodium and chlorine atoms and their ions in a coupled channel treatment†

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In an extension of previous work (Šimsová *et al.*, *Phys. Chem. Chem. Phys.*, 2022, **24**, 25250), we study stimulated radiative association of sodium chloride (NaCl) in an environment with a black body radiation. Colliding neutral (Na and Cl) and ionic (Na⁺ and Cl⁻) fragments are considered. The coupling between the diabatic ionic and neutral channels is accounted for. The cross sections are computed and resolved on the vibrational states of the formed NaCl molecule for detailed analysis. The thermal rate coefficients for neutral colliding fragments at kinetic temperatures, *T*, from 1 K to 5300 K are computed for use in astrochemical modelling. The total rate coefficient is affected by more than one order of magnitude by stimulated emission from a blackbody radiator of temperature *T*_b = 50 000 K. The effect from stimulated emission is largest for the lowest kinetic temperatures, where *T*_b of a few thousand kelvins has a significant effect. The rate coefficient for the colliding ionic fragments is calculated from 80 K to 3615 K. The blackbody radiation has little effect on this process.

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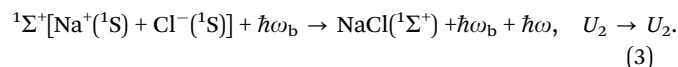
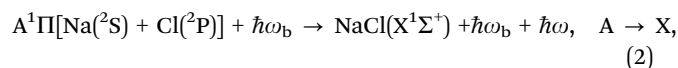
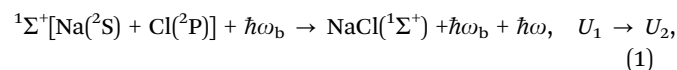
1 Introduction

Expanding circumstellar nebulae are an important part of the interstellar medium. They are formed by materials ejected from stars, and they essentially work as giant space chemical laboratories, where molecules and dust are formed and destroyed. Their environment can be very dynamic with shocks, winds and high-speed outflows, and there is the stellar and interstellar radiation always being present; even when the stellar radiation is absorbed by an optically thick circumstellar dust shell it is often re-radiated at longer wavelengths.

Sodium and potassium chlorides (NaCl and KCl) have been detected several times in the interstellar medium. They were observed in the nebulae surrounding carbon-rich stars CW Leonis^{1,2} and V1610 Cygnus³ and also oxygen-rich stars VY Canis Majoris,⁴ IK Tauri,⁵ and QX Puppis,⁶ where the corresponding central objects are stars in their asymptotic giant branch⁷ (AGB) or post-AGB⁸ evolution stages. However, fairly

recently, NaCl was also observed in the protostellar disk around the young accreting high-mass star Orion Src I.⁹

The process of radiative association belongs to the important chemical processes responsible for molecule formations in the circumstellar nebulae.¹⁰ While the experimental investigations of radiative association are very difficult,¹¹ the theoretical studies have been plentiful in the last decades.^{12–25} Little is known about the formation of NaCl.²⁶ In our previous work, making use of the detailed diabatic picture with coupling, we studied the spontaneous radiative association to form NaCl^{27–29} and also the non-radiative charge transfer, *i.e.*, chemiionisation and mutual neutralisation.²⁹ However, the ever-present radiation field provides a permanent source of radiation which by stimulating the radiative processes could significantly enhance the corresponding rates.¹⁰ For example, in the case of radiative association the rate coefficients could be increased even by factors of 100 to 1000 in the low temperature regions.^{30–38} That is why, in this work, we study the stimulated radiative association of NaCl, specifically the processes



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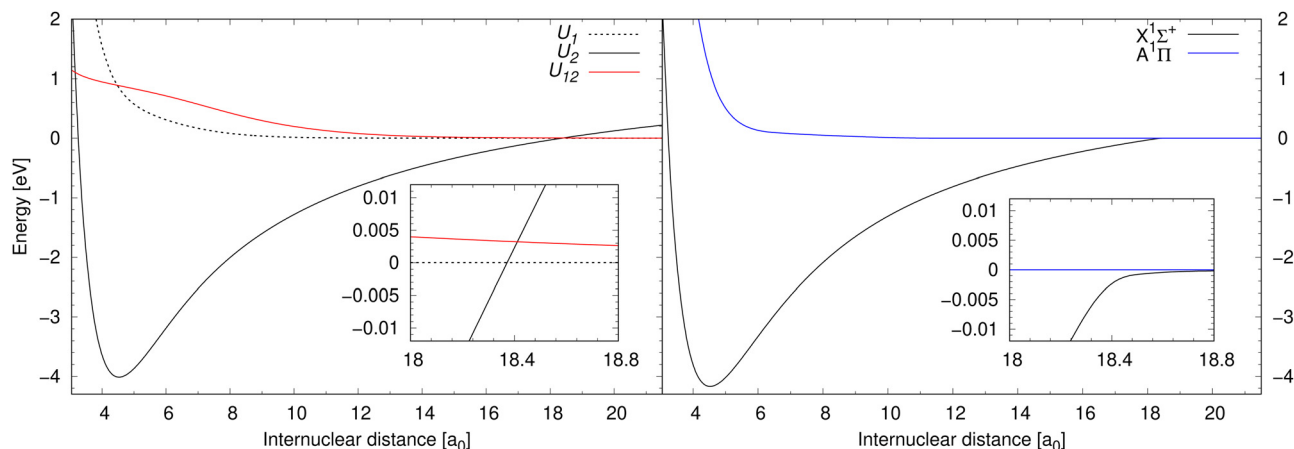


Fig. 1 Left panel: Potential energy curves for the electronic states U_1 and U_2 of $^1\Sigma^+$ symmetry in the diabatic representation between which processes (1) and (3) are calculated are illustrated together with their diabatic coupling U_{12} . Right panel: Illustrates potential energy curves for the electronic states A and X of $^1\Pi$ and $^1\Sigma^+$ symmetry, respectively, in the adiabatic representation, between which process (2) is calculated. The data for the $^1\Sigma^+$ symmetry are taken from Giese & York,³⁹ and the data for the $A^1\Pi$ state are taken from Zeiri & Balint-Kurti.⁴⁰

Process (1) and process (2) are collisional processes of the same reactants (neutral sodium and chlorine atoms in their atomic ground electronic states), however, in different molecular symmetries. While the molecular state of the reactants in process (1) is in the $^1\Sigma^+$ symmetry, the molecular state of the reactants in process (2) is in the $^1\Pi$ symmetry. As process (1) here is treated in the diabatic representation (in the coupled channel treatment), we will simply denote the initial and final electronic states as U_1 and U_2 . Since process (2) is treated conventionally (uncoupled solution of the radial nuclear Schrödinger equation), we will denote the initial and final electronic states as is customary in the Born–Oppenheimer approximation: A and X. Finally, process (3) is a collisional process

of a sodium cation and a chloride anion in their atomic ground electronic states. The molecular state of the reactants in process (3) is in the $^1\Sigma^+$ symmetry, so in the diabatic representation the initial and final electronic state is U_2 . The products of processes (1)–(3) are the same, concretely sodium chloride in the lowest electronic state of the $^1\Sigma^+$ symmetry. Notice that during process (2), the molecular electronic angular momentum is changed. Potential energy curves in the diabatic representation of processes (1) and (3) are illustrated in the left-hand panel of Fig. 1 together with their diabatic coupling. Potential energy curves in the adiabatic representation (in the Born–Oppenheimer approximation) of process (2) are illustrated in the right-hand panel of Fig. 1.

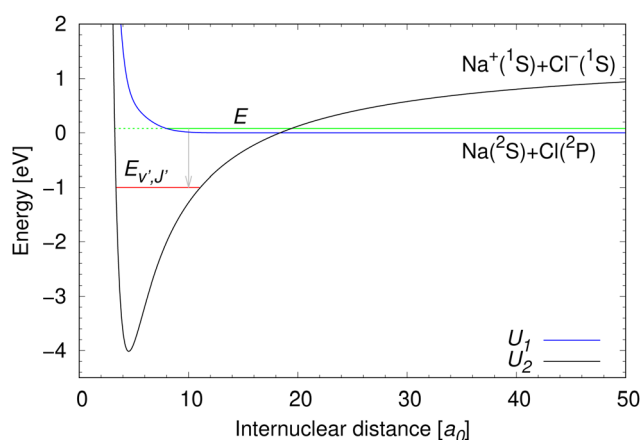


Fig. 2 Potential energy curves of the two lowest $^1\Sigma^+$ diabatic electronic states of NaCl.³⁹ The meaning of collision energy E , energy of final bound ro-vibrational states $E_{v,J}$, and the dissociation limits of the diabatic $U_1(R)$ and $U_2(R)$ potential energy curves are illustrated. The initial state lives on both diabatic electronic states, which is indicated by the green full and dashed lines.

2 Methods

The same *ab initio* data^{39,40} are used as in our previous work on NaCl.^{27–29} The potential energy curves for the two lowest $^1\Sigma^+$ electronic states of NaCl in the diabatic representation are illustrated in Fig. 2.

2.1 Rate coefficients and cross sections

The rate coefficients for formation of a molecule with the reduced mass μ by stimulated radiative association at temperature T is determined by

$$\alpha(T; T_b) = \left(\frac{8}{\mu\pi}\right)^{1/2} \left(\frac{1}{k_B T}\right)^{3/2} \int_0^\infty E\sigma(E; T_b)e^{-E/k_B T} dE, \quad (4)$$

where k_B is the Boltzmann constant and the temperature T_b characterises the black-body radiation field and

$$\sigma(E; T_b) = \sum_{J;v',J'} \sigma_{J;v',J'}(E; T_b). \quad (5)$$

In the case of dipole moment transitions, the individual cross sections of stimulated radiative association can be expressed as³⁰

$$\sigma_{J;v',J'}(E; T_b) = \frac{1}{4\pi\epsilon_0} \frac{8\pi^2}{3} \frac{\omega_{J;v',J'}^3}{k^2} p \left(\frac{\omega_{J;v',J'}}{c} \right)^3 |H_{J \rightarrow J'}| |M_{J;v',J'}(E)|^2 \times \left[\frac{1}{1 - \exp(-\hbar\omega_{J;v',J'}/k_B T_b)} \right], \quad (6)$$

where ϵ_0 denotes the vacuum permittivity, J, J' denote the rotational quantum numbers of the initial and final states, respectively, v' is the vibrational quantum number of the final bound state, $k = \sqrt{2\mu E}/\hbar$, c is the speed of light and $H_{J \rightarrow J'}$ are the Hönl-London factors (obviously, only final states v', J' , for which $H_{J \rightarrow J'} \neq 0$ are eligible for radiative association from the initial state E, J). The probabilities p of approach in the initial channel are 1/12 for process (1), 1/6 for process (2), and 1 for process (3). $\omega_{J;v',J'}(E)$ is the emitted photon angular frequency, $\hbar\omega_{J;v',J'}(E) = E - E_{v',J'}$, where $E > 0$ and $E_{v',J'} < 0$ are taken with respect to the neutral dissociation limit. $M_{J;v',J'}(E)$ is the dipole-moment matrix element between the wavefunction of the initial continuum state and the wavefunction of the final bound state. The corresponding wavefunctions are obtained in the same manner as in our studies of spontaneous radiative association,^{27–29} *i.e.* within the coupled diabatic picture for processes (1) and (3), and within the adiabatic approximation for process (2).

2.2 Numerical details

The cross sections for radiative association for processes (1), (2), and (3) were calculated from 10^{-5} eV up to 10 eV for background temperatures: 50 K, 100 K, 150 K, 200 K, 500 K, 750 K, 1000 K, 2000 K, 3000 K, 4000 K, 5000 K, 7500 K, 10 000 K, 25 000 K, and 50 000 K. The spontaneous emission case ($T_b = 0$ K) was calculated in previous studies.^{27–29} Similarly as in those studies, processes (1) and (3) were treated by non-adiabatic dynamics, while process (2) was treated by conventional perturbation theory. For process (1), the same input parameters (minimum and maximum internuclear distances for both the bound and continuum wavefunction calculations, including the radial step, and maximum rotational quantum number) as previously^{27–29} were used for collision energies from 0.03 eV and up, for all background temperatures. Only at low collision energies, where the effect of background temperature is profound, did we increase the included rotational quantum numbers and consequently also the maximum internuclear distance for the Numerov method. Concretely, for process (1), we set $J_{\max} = 20$ and $R_{\max} = 45a_0$ up to 10^{-4} eV, $J_{\max} = 23, 30, 40$ up to 0.0004 eV, 0.0007 eV, and 0.001 eV, respectively, with $R_{\max} = 23a_0$. $R_{\max} = 23a_0$ was used up to 0.03 eV with the same J_{\max} as by Gustafsson.²⁷ Lastly, the same energy grid as before^{27–29} suffices for rate-coefficient convergence. For process (2), the convergence was achieved by the same input parameters as previously.^{27–29} For process (3), much lower maximum internuclear distance, $23a_0$ up to 4.6 eV and $30a_0$ above, was used here in comparison with Šimsová née Zámečníková *et al.*²⁹ as there we simultaneously calculated non-radiative continuum-continuum processes for which large internuclear

distances are needed in solving the radial Schrödinger equation.

3 Results and discussion

Cross sections are illustrated in Fig. 3 for eight non-zero background temperatures and compared to the spontaneous emission. The recalculated cross section for $T_b = 0$ K coincides with the spontaneous emission obtained by Gustafsson,²⁷ Šimsová-Zámečníková *et al.*²⁸ and Šimsová née Zámečníková *et al.*²⁹ Typically, the higher the background temperature, the larger the cross sections become.^{30–32,35,36} Also, the effect is most noticeable at low collision energies and becomes insignificant at large collision energies.

As the energy surpasses the opening of the ionic scattering channel, at 1.4845 eV, the resonances for process (1) cease. This can be understood since there are (by definition) no bound states on the ionic potential, U_2 , at those energies. Bound states on the U_2 potential are those that give rise to Fano-Feshbach resonances when the non-adiabatic coupling is accounted for. Furthermore, no shape (orbiting) resonances are supported above the 1.4845 eV threshold either. Shape resonances would

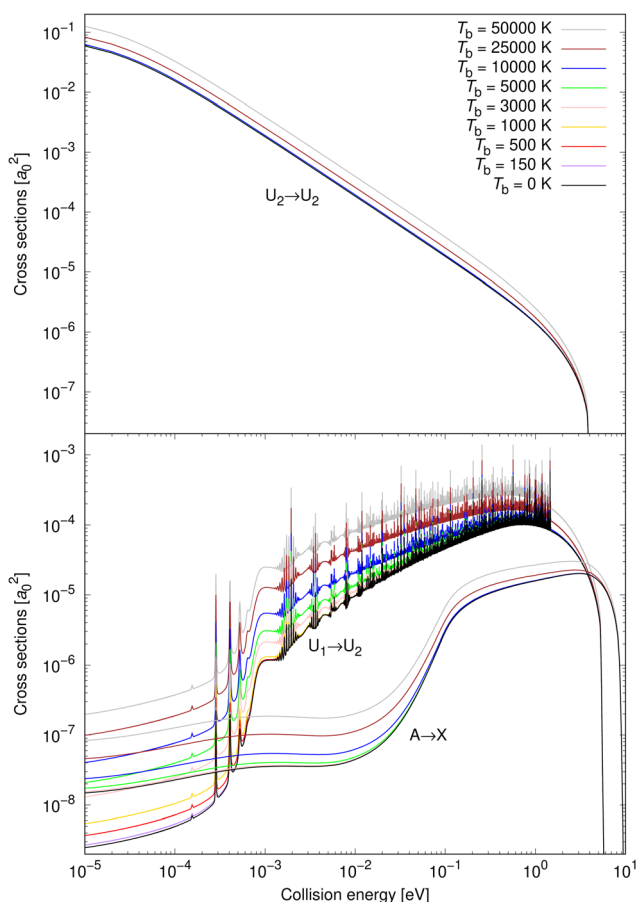


Fig. 3 Radiative association cross sections for several background temperatures, including zero. Top: Calculated by the non-adiabatic dynamics for process (3); bottom: calculated by the non-adiabatic dynamics for process (1) and conventionally for process (2).

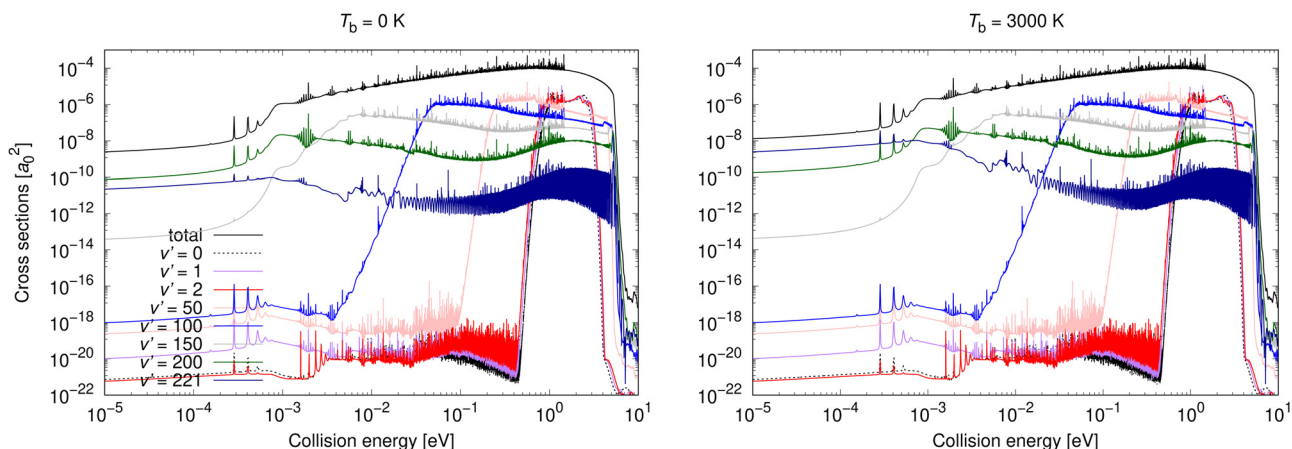


Fig. 4 Radiative association vibrationally resolved cross sections obtained allowing for non-adiabatic dynamics for process (1) for two background temperatures: 0 K and 3000 K.

appear if the effective ionic potential had a barrier peaking above 1.4845 eV, which would permit quasibound states. There is, however, no such barrier. The reason is that the ionic potential is attractive with a $1/R$ dependence while the repulsive centrifugal potential has a $1/R^2$ dependence. Since the attractive term falls off more slowly than the repulsive, a centrifugal barrier cannot form.

In the top panel of Fig. 3, the cross sections for process (3) are illustrated for the same set of background temperatures as for process (1) and process (2). The background enhances the radiative association and is most easily seen for the higher background temperatures. The enhancement becomes less significant as the collision energy increases.

Vibrationally resolved cross sections were calculated for process (1) for chosen vibrational quantum levels of the target electronic state: $v' = 0, 1, 2, 25, 50, 75, 100, 125, 150, 175, 200$, and 221, with the last one being the highest vibrational level of the U_2 diabatic state below the energetic asymptotic limit of the neutral atoms. These calculations were performed for all studied background temperatures and for the spontaneous emission. In Fig. 4, some of these vibrationally resolved cross

sections are compared for $T_b = 0$ and 3000 K. The total cross sections for these processes are also illustrated in Fig. 4.

From Fig. 4 it can be seen that at low collision energies the lowest vibrational levels do not contribute much to the cross sections for either of the shown background temperatures. The order of the importance of the vibrationally resolved cross sections depends on the background temperature. In particular, at low collision energies, transitions to $v' = 221$ are less important at $T_b = 0$ K than at $T_b = 3000$ K, where this level contributes most of all levels to the total cross section. With increasing collision energy, the importance of lower vibrational levels increases. Typically many levels contribute significantly to the total cross section. The vibrationally resolved cross sections at all studied background temperatures are illustrated in the ESI.†

Vibrationally resolved cross sections were calculated also for process (2) for almost the same chosen vibrational levels of the target electronic state: 0, 1, 2, 25, 50, 75, 100, 125, 150, 175, 200. Some of them are illustrated for $T_b = 0$ and 10 000 K in Fig. 5.

Neither transitions to the lowest, nor to the highest, vibrational levels contribute significantly at low collision energies,

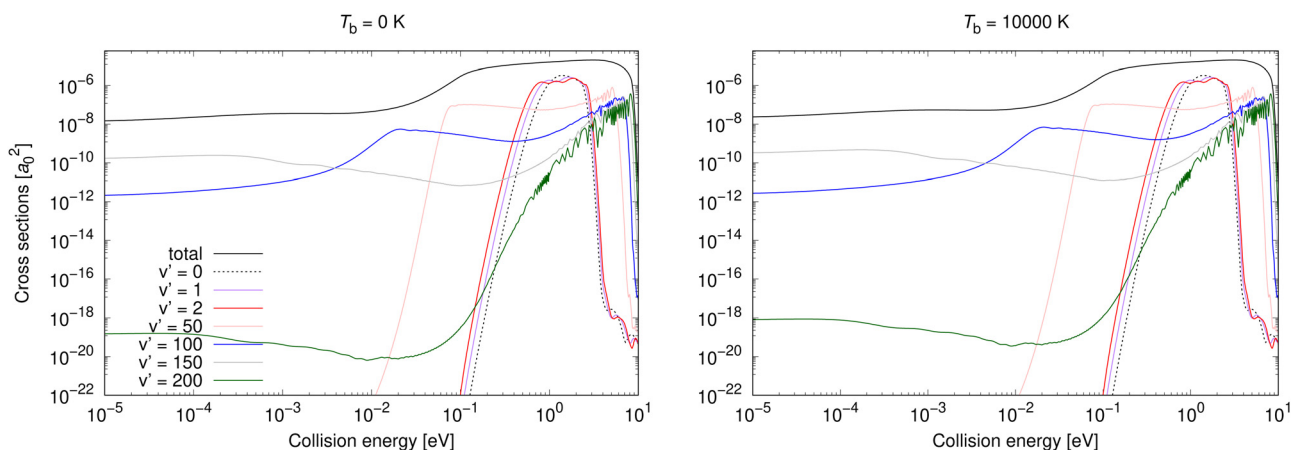


Fig. 5 Radiative association vibrationally resolved cross sections for process (2) for two background temperatures, 0 K and 10 000 K.

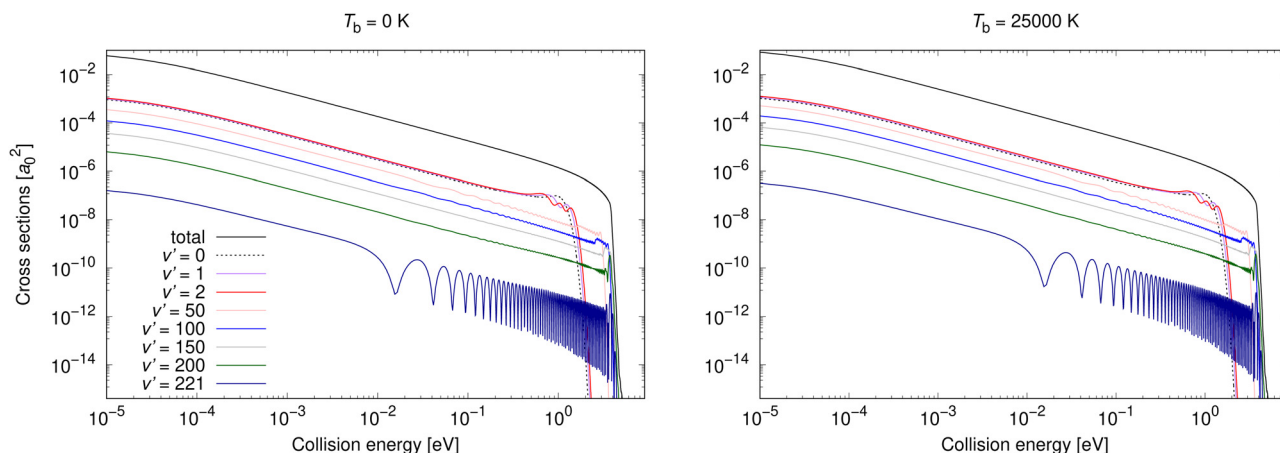


Fig. 6 Radiative association vibrationally resolved cross sections for process (3) calculated with non-adiabatic dynamics for two background temperatures, 0 K and 25 000 K.

which is valid for both shown background temperatures. This is different from process (1), where the highest vibrational level contributes more at both background temperatures, and the

most at $T_b = 3000$ K. At higher collision energies, transitions to all vibrational levels contribute more similarly to the total cross section. The vibrationally resolved cross sections at all studied background temperatures are illustrated in the ESI.†

Gustafsson²⁷ and Šimsová-Zámečníková *et al.*²⁸ showed that the uncoupled solution for radiative association of $\text{Na}(^2\text{S})$ and $\text{Cl}(^2\text{P})$ in the Σ^+ -symmetry vastly underestimates the correct cross sections. Therefore, we did not study stimulated emission in that way here. However, we did calculate vibrationally resolved cross sections, which can be compared to the non-adiabatically obtained ones. These results are in the ESI,† for energies up to 0.03 eV.

Non-adiabatic vibrationally resolved cross sections for $T_b = 0$ K and 25 000 K for process (3) are illustrated in Fig. 6. As expected from the total cross sections from the top panel of Fig. 3, the vibrationally resolved cross sections for these background temperatures do not differ much between these background temperatures. The vibrationally resolved cross sections at other studied background temperatures are shown in the ESI,† where it can be seen that the order of importance of vibrational levels does not depend on the background temperature.

Rate coefficients for processes (1) and (2) were calculated from 1 K to 5300 K for all studied background temperatures. They are illustrated for nine background temperatures in the bottom panel of Fig. 7. Except for the very lowest temperatures, process (1) dominates over process (2), for all background temperatures. At background temperatures above about 1000 K, process (2) starts to be less efficient even at the lowest temperatures. Similarly as for cross sections, it can be seen that a non-zero background temperature enhances the rate coefficients; and the larger the background temperature is, the larger the enhancement is. The relative importance of a non-zero background temperature drops at larger temperatures.^{30–32,35–38}

The total rate coefficient for formation of NaCl by radiative association of sodium and chlorine in their ground electronic states at several background temperatures were computed by summing the rate coefficients of processes (1) and (2). They are illustrated at chosen background temperatures in the top-panel

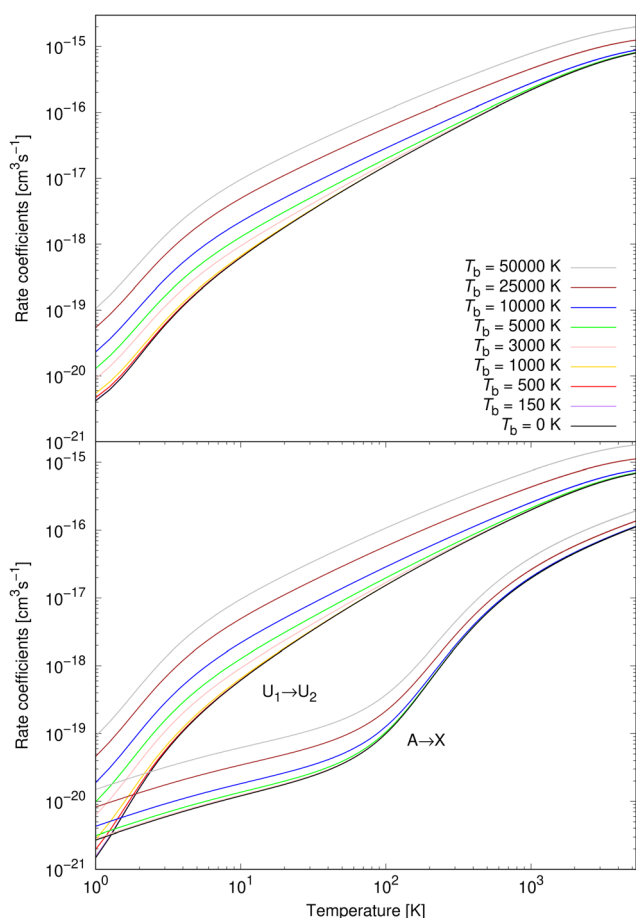


Fig. 7 Top panel: Total radiative association rate coefficients summed over processes (2) and (1), at several background temperatures. Bottom panel: Radiative association rate coefficients for processes (2) and (1), at several background temperatures.

Table 1 Kooij-function fitting parameters are summarised for the total radiative association of Na(²S) + Cl(²P) and radiative association of Na⁺(¹S) + Cl⁻(¹S) at $T_b = 10\,000$ K, 25 000 K, and 50 000 K

Process	T [K]	A [cm ³ s ⁻¹]	B	C [K]
$T_b = 10\,000$ K				
RA Na + Cl	1–1.675	1.9026(–12)	3.48386	–1.64608
	1.675–2.3	1.5998(–14)	2.34032	0.29014
	2.3–3.425	2.7916(–16)	1.29542	2.68895
	3.425–11.25	6.8600(–17)	0.89322	4.04831
	11.25–30	8.9452(–17)	1.01184	2.63814
	30–80	9.0080(–17)	1.00878	3.05820
	80–350	8.8920(–17)	0.97253	5.90601
	350–750	9.7141(–17)	0.90087	33.2892
	750–1400	1.2993(–16)	0.75034	148.868
	1400–2400	2.3008(–16)	0.52634	467.489
	2400–4000	5.3958(–16)	0.25019	1137.98
	4000–5300	1.2069(–15)	2.3396(–2)	2010.80
	RA Na ⁺ + Cl ⁻	80–500	1.4224(–16)	–0.50804
500–1600		1.5036(–16)	–0.54326	21.1331
1600–3615		1.8809(–16)	–0.62599	159.760
$T_b = 25\,000$ K				
RA Na + Cl	1–1.675	3.4001(–12)	3.41917	–1.54738
	1.675–2.3	3.1304(–14)	2.29762	0.34958
	2.3–3.425	5.8098(–16)	1.26849	2.71361
	3.425–11.25	1.3582(–16)	0.85108	4.13555
	11.25–30	1.7021(–16)	0.95308	2.89686
	30–80	1.6985(–16)	0.94287	3.54081
	80–301	1.6724(–16)	0.88484	8.54798
	301–740	1.8250(–16)	0.80267	35.2174
	740–1400	2.4400(–16)	0.64995	149.129
	1400–2400	4.1951(–16)	0.43740	451.090
	2400–4000	9.6469(–16)	0.16778	1106.95
	4000–5300	2.1678(–15)	–6.0313(–2)	1985.02
	RA Na ⁺ + Cl ⁻	80–480	1.8903(–16)	–0.51142
480–1600		2.0263(–16)	–0.55634	26.0469
1600–3615		2.6370(–16)	–0.65416	187.919
$T_b = 50\,000$ K				
RA Na + Cl	1–1.675	6.0463(–12)	3.39642	–1.51413
	1.675–2.3	5.7719(–14)	2.28361	0.36746
	2.3–3.425	1.0930(–15)	1.25964	2.72011
	3.425–11.25	2.5080(–16)	0.83655	4.16533
	11.25–30	3.0949(–16)	0.93202	2.99521
	30–80	3.0744(–16)	0.91826	3.75081
	80–268	3.0215(–16)	0.85341	9.32976
	268–676	3.2717(–16)	0.76856	33.6494
	676–1400	4.3781(–16)	0.60936	144.302
	1400–2400	7.6993(–16)	0.38673	456.484
	2400–4000	1.7748(–15)	0.11639	1114.63
	4000–5300	4.0480(–15)	–0.11588	2009.00
	RA Na ⁺ + Cl ⁻	80–460	2.8797(–16)	–0.51330
460–1499		3.0963(–16)	–0.56171	26.8042
1499–2000		3.6014(–16)	–0.62258	107.663
2000–3615		4.4126(–16)	–0.69081	255.751

$$x(-y) \equiv x \times 10^{-y}.$$

of Fig. 7. Additionally, the total rate coefficients at several background temperatures were fitted to the Kooij function of the Arrhenius type

$$f(T) = A \left(\frac{T}{300} \right)^B \exp \left(-\frac{C}{T} \right). \quad (7)$$

The fitting parameters A , B , C are summarised in Table 1 for $T_b = 10\,000$ K, 25 000 K, and 50 000 K. The fitting parameters for

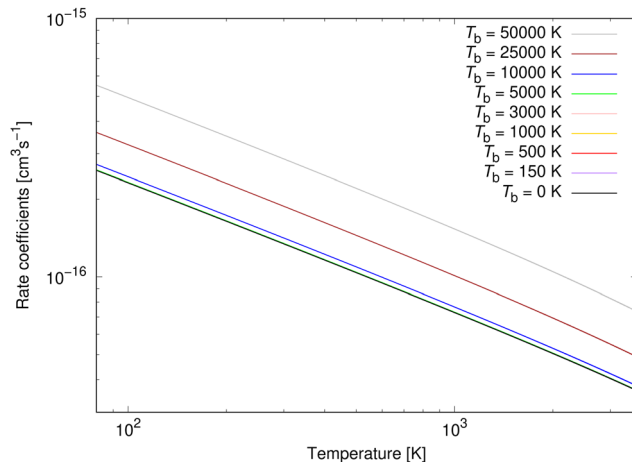


Fig. 8 Radiative association rate coefficients for processes (3) at several background temperatures.

the total rate coefficients at background temperatures from 500 K to 7500 K are summarised in the ESL.†

The rate coefficients for process (3) were calculated up to 3615 K as in Šimsová née Zámečníková *et al.*²⁹ They are illustrated for several chosen temperatures in Fig. 8. As expected from the cross sections in the top panel of Fig. 3, the background radiation has little effect on the rate coefficients of process (3). The rate coefficients up to $T_b = 5000$ K coincide with the case where no blackbody radiation is present. Because of the lack of the effect, we summarise the fitting parameters to the Kooij function (7) for only $T_b = 10\,000$ K, 25 000 K, and 50 000 K.

4 Conclusions

We have computed stimulated radiative association cross sections and rate coefficients for the formation of NaCl molecules in their electronic ground state. We also present vibrationally resolved cross sections, *i.e.* for selected vibrational states of the formed molecule. The emission is stimulated by a blackbody radiation of temperature, T_b , which we vary from 0 to 50 000 K. Overall the stimulated emission plays the biggest role at low collision energies, and at low kinetic temperatures, T .

Two reactions for neutral colliding fragments, Na and Cl, are considered. One of them, process (1), involves non-adiabatic couplings and its cross section largely overshadows that of the other, process (2), except for at the lowest and the highest energies. For values of T_b of a few hundred kelvins, the cross section is only affected at energies below ~ 1 meV. This is the energy above which non-adiabatic dynamics starts to play a big role,²⁷ meaning that the emitted photons are higher in energy above 1 meV of collision energy. This can be seen in the vibrationally resolved cross section for process (1), where transitions to the highest vibrational states dominate below 1 meV. As can be seen in eqn (6), for stimulated emission to make a difference $k_b T_b$ should be comparable to, or bigger than, the typical energy of the emitted photons.

The cross section for process (1) has pronounced resonances all the way up to a collision energy of 1.4845 eV, which is the threshold energy for the ionic channel. Due to the $1/R$ behaviour of the ionic state, there are no resonances above that energy (see Section 3). The other neutral reaction, process (2), has no resonances since the potential of the entrance state, A, is purely repulsive. Shape resonances, if present, have been observed to contribute to rate coefficients mainly at low temperatures.¹⁸ In the Born–Oppenheimer approximation there are only shape resonances. In the diabatic representation there are also Fano–Feshbach resonances. While shape resonances have only additive contributions to cross sections, the Fano–Feshbach resonances can have both additive and subtractive effects on cross sections and hence also on rate coefficients.

The third reaction, process (3), which corresponds to colliding ions, Na^+ and Cl^- , is also treated with non-adiabatic couplings. It is completely resonance free, since the entrance channel is open only above the ionic threshold. The same argument as for process (1) holds here. The corresponding rate coefficient falls off monotonically with increasing kinetic temperature, T , and it is significantly affected by stimulated emission only if the radiation temperature is 10 000 K or more.

The total rate coefficient for colliding neutral atoms, Na and Cl, is affected by stimulated emission up to two orders of magnitude at the lowest kinetic temperatures, T , for the highest radiation temperatures, T_b . This is expected to be important in astrochemical modelling of environments with abundant radiation.

Conflicts of interest

There are no conflicts to declare.

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Notes and references

- J. Cernicharo and M. Guélin, *Astron. Astrophys.*, 1987, **183**, L10–L12.
- M. Agúndez, J. P. Fonfría, J. Cernicharo, C. Kahane, F. Daniel and M. Guélin, *Astron. Astrophys.*, 2012, **543**, A48.
- J. L. Highberger, K. J. Thomson, P. A. Young, D. Arnett and L. M. Ziurys, *Astrophys. J.*, 2003, **593**, 393–401.
- L. M. Ziurys, S. N. Milam, A. J. Apponi and N. J. Woolf, *Nature*, 2007, **447**, 1094–1097.
- S. N. Milam, A. J. Apponi, N. J. Woolf and L. M. Ziurys, *Astrophys. J.*, 2007, **668**, L131–L134.
- C. Sánchez Contreras, J. Alcolea, V. Bujarrabal and A. Castro-Carrizo, *Astron. Astrophys.*, 2018, **618**, A164.
- F. Herwig, *Ann. Revue Astron. Astrophys.*, 2005, **43**, 435–479.
- H. Van Winckel, *Ann. Revue Astron. Astrophys.*, 2003, **41**, 391–427.
- A. Ginsburg, B. McGuire, R. Plambeck, J. Bally, C. Goddi and M. Wright, *Astrophys. J.*, 2019, **872**, 54.
- A. E. Glassgold, *Ann. Revue Astron. Astrophys.*, 1996, **34**, 241–277.
- D. Gerlich and S. Horning, *Chem. Rev.*, 1992, **92**, 1509–1539.
- D. R. Bates, *Mon. Not. R. Astron. Soc.*, 1951, **111**, 303–314.
- J. A. Babb and A. Dalgarno, *Phys. Rev. A*, 1995, **51**, 3021–3026.
- F. A. Gianturco and P. G. Giorgi, *Astrophys. J.*, 1997, **479**, 560–567.
- A. S. Dickinson and F. X. Gadéa, *Mon. Not. R. Astron. Soc.*, 2000, **318**, 1227–1231.
- F. Mrugała, V. Špirko and W. P. Kraemer, *J. Chem. Phys.*, 2003, **118**, 10547–10560.
- X. J. Liu, Y. Z. Qu, B. J. Xiao, C. H. Liu, Y. Zhou, J. G. Wang and R. J. Buenker, *Phys. Rev. A*, 2010, **81**, 022717.
- L. Augustovičová, V. Špirko, W. P. Kraemer and P. Soldán, *Chem. Phys. Lett.*, 2012, **531**, 59–63.
- T. Stoecklin, F. Lique and M. Hochlaf, *Phys. Chem. Chem. Phys.*, 2013, **15**, 13818–13825.
- G. Nyman, M. Gustafsson and S. V. Antipov, *Int. Rev. Phys. Chem.*, 2015, **34**, 385–428.
- M. Zámečnicková, P. Soldán, M. Gustafsson and G. Nyman, *Mon. Not. R. Astron. Soc.*, 2019, **489**, 2954–2960.
- M. Zámečnicková, M. Gustafsson, G. Nyman and P. Soldán, *Mon. Not. R. Astron. Soc.*, 2020, **492**, 3794–3802.
- D. Jones (née Burdakova), M. Gustafsson and G. Nyman, *Mon. Not. R. Astron. Soc.*, 2022, **517**, 4892–4901.
- T. Bai, Z. Qin and L. Liu, *Mon. Not. R. Astron. Soc.*, 2022, **510**, 1649–1656.
- P. Szabó and M. Gustafsson, *J. Chem. Phys.*, 2023, **159**, 144112.
- T. J. Millar, C. Walsh, M. Van de Sande and A. J. Markwick, *Astron. Astrophys.*, 2023, DOI: [10.1051/0004-6361/202346908](https://doi.org/10.1051/0004-6361/202346908).
- M. Gustafsson, *J. Chem. Phys.*, 2020, **153**, 114305.
- M. Šimsová-Zámečnicková, P. Soldán and M. Gustafsson, *Astron. Astrophys.*, 2022, **664**, A5.
- M. Šimsová née Zámečnicková, M. Gustafsson and P. Soldán, *Phys. Chem. Chem. Phys.*, 2022, **24**, 25250–25257.
- P. C. Stancil and A. Dalgarno, *Astrophys. J.*, 1997, **479**, 543–546.
- P. C. Stancil and A. Dalgarno, *Astrophys. J.*, 1997, **490**, 76–78.
- B. Zygelman, P. C. Stancil and A. Dalgarno, *Astrophys. J.*, 1998, **508**, 151–156.
- A. S. Dickinson, *J. Phys. B: At., Mol. Opt. Phys.*, 2005, **38**, 4329–4334; A. S. Dickinson, *J. Phys. B: At., Mol. Opt. Phys.*, 2008, **41**, 049801.

- 34 S. Bovino, M. Tacconi and F. A. Gianturco, *Astrophys. J.*, 2011, **740**, 101; S. Bovino, M. Tacconi and F. A. Gianturco, *Astrophys. J.*, 2012, **748**, 150.
- 35 L. Augustovičová, V. Špirko, W. P. Kraemer and P. Soldán, *Mon. Not. R. Astron. Soc.*, 2013, **435**, 1541–1546.
- 36 L. Augustovičová, W. P. Kraemer and P. Soldán, *Astrophys. J.*, 2014, **782**, 46.
- 37 L. Augustovičová, M. Zámečnicková, W. P. Kraemer and P. Soldán, *Chem. Phys.*, 2015, **462**, 65–70.
- 38 M. Zámečnicková, W. P. Kraemer and P. Soldán, *J. Quant. Spectrosc. Radiat. Transfer*, 2017, **191**, 88–95.
- 39 T. J. Giese and D. M. York, *J. Chem. Phys.*, 2004, **120**, 7939–7948.
- 40 Y. Zeiri and G. G. Balint-Kurti, *J. Chem. Phys.*, 1983, **99**, 1–24.